

An imaginary QPO in the soft-to-hard transition of MAXI J1820+070

Candela A. Bellavita

Instituto Argentino de Radioastronomía (IAR/CONICET)



CONICET



Instituto Argentino
de Radioastronomía



kapteyn astronomical
institute

Mariano Méndez (RUG, Neetherlands)
Federico García (IAR/CONICET)
Ruican Ma (CAS, China)
Ole König (CfA,USA; FAU, Germany)
Jörn Wilms (FAU, Germany)

July 2024

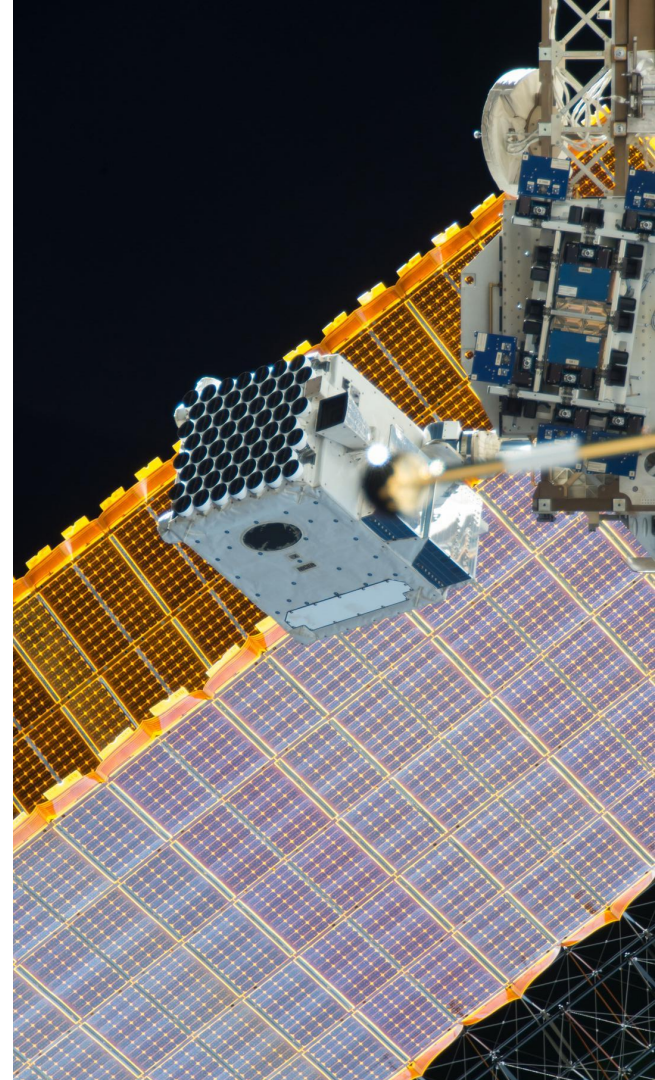
MAXI J1820+070

BH X-ray binary

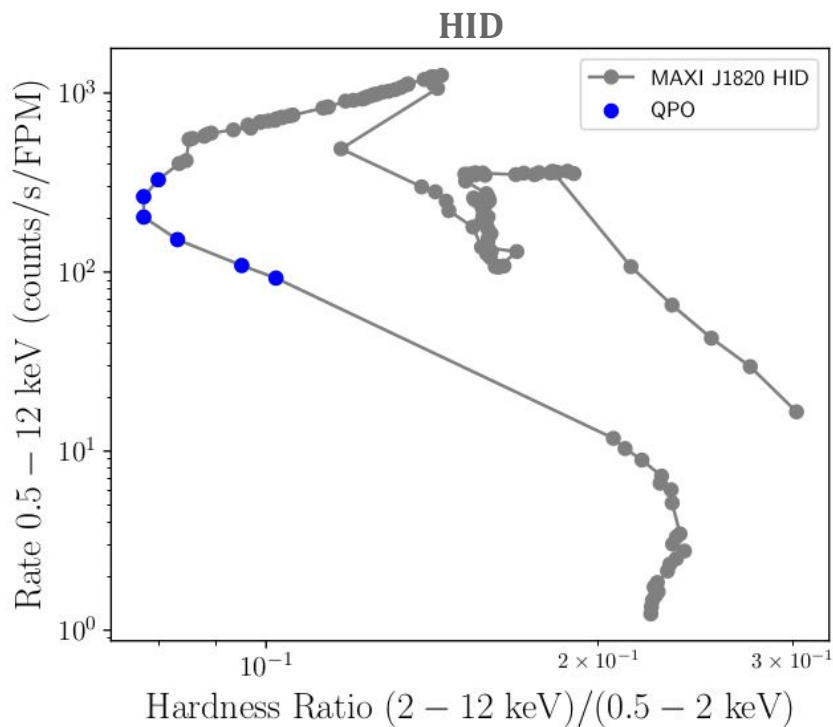
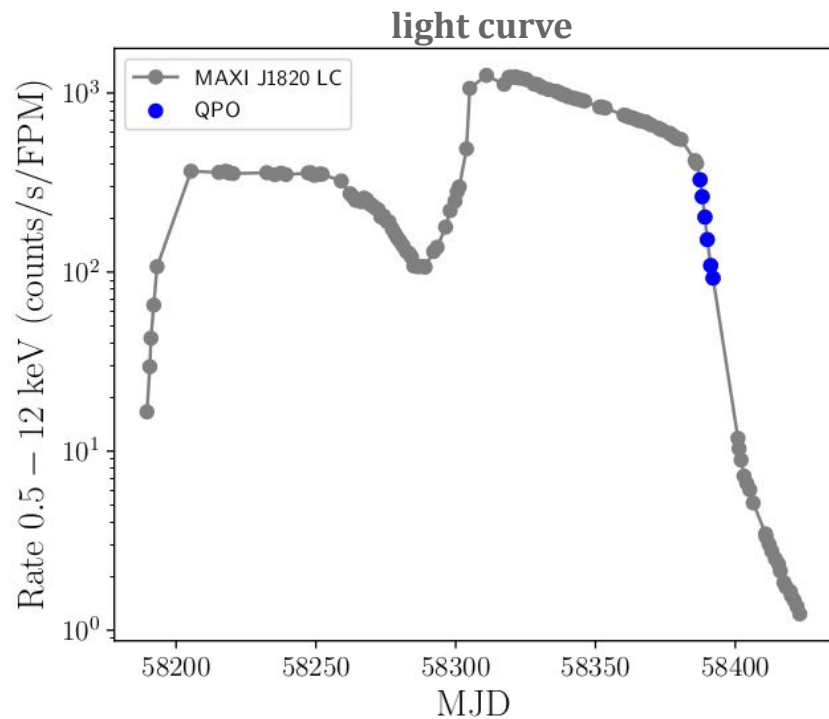
underwent an **outburst** in 2018



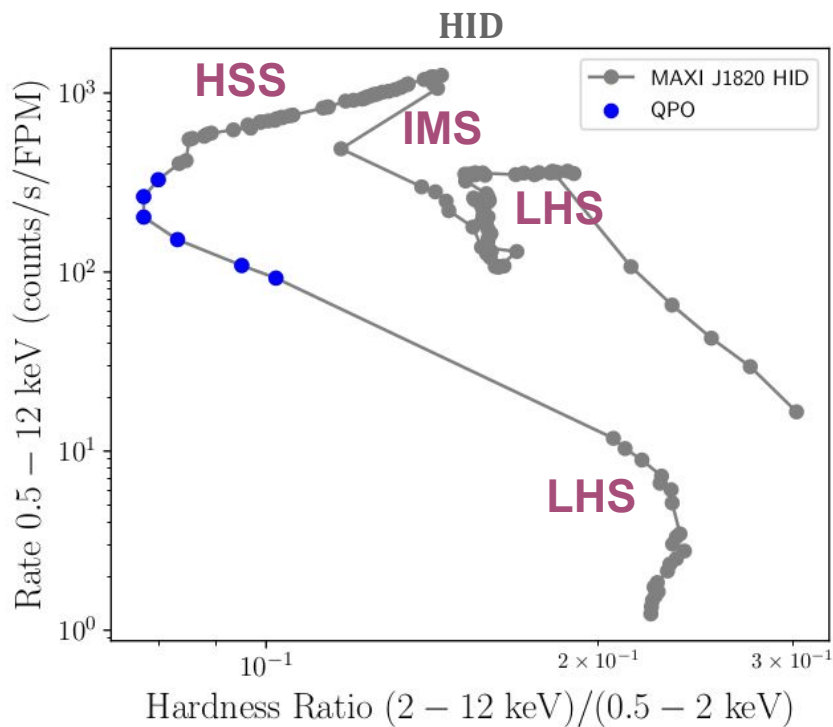
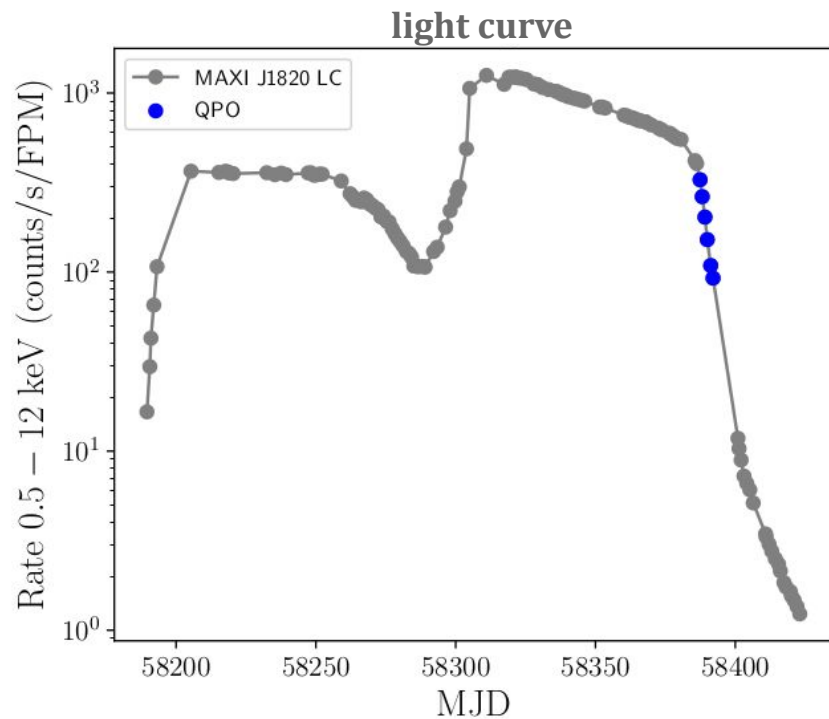
closely monitored by **NICER**



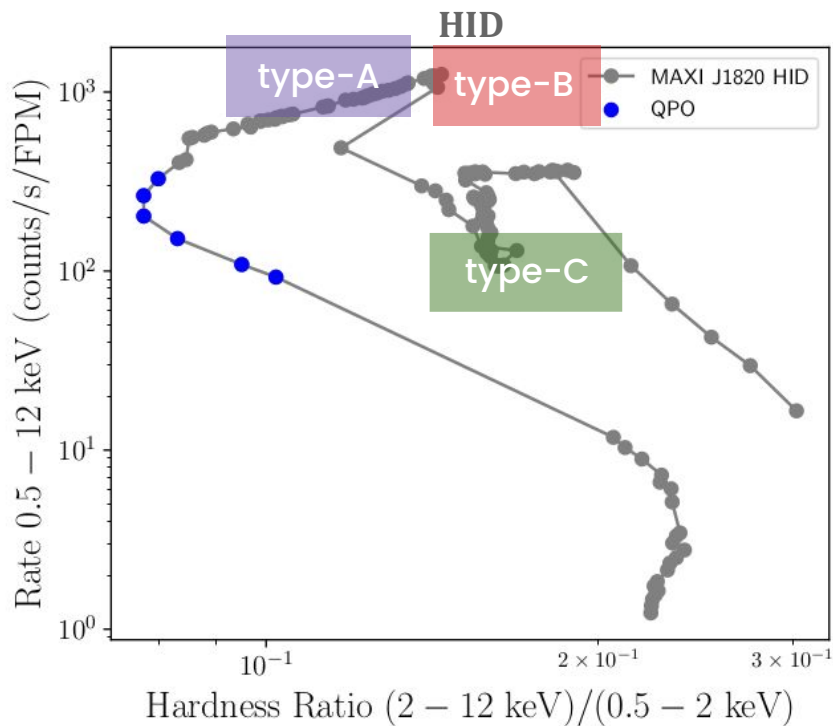
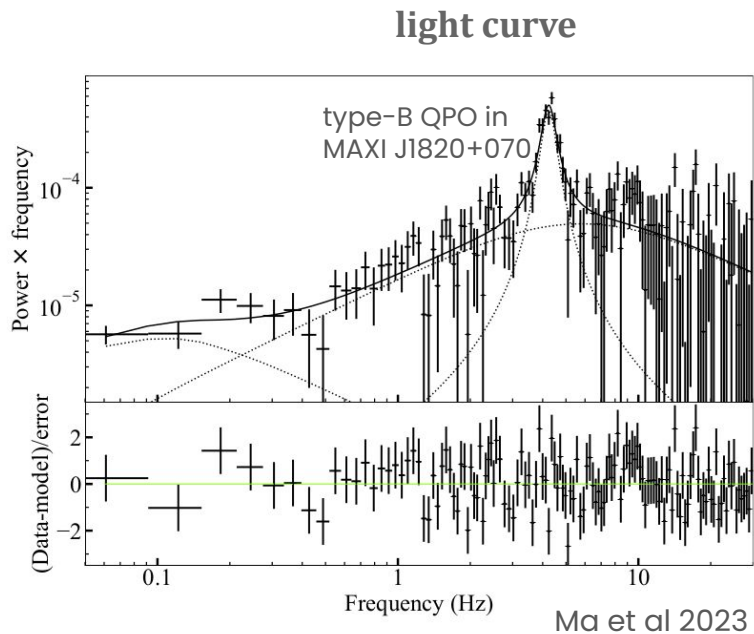
MAXI J1820+070: light curve & HID



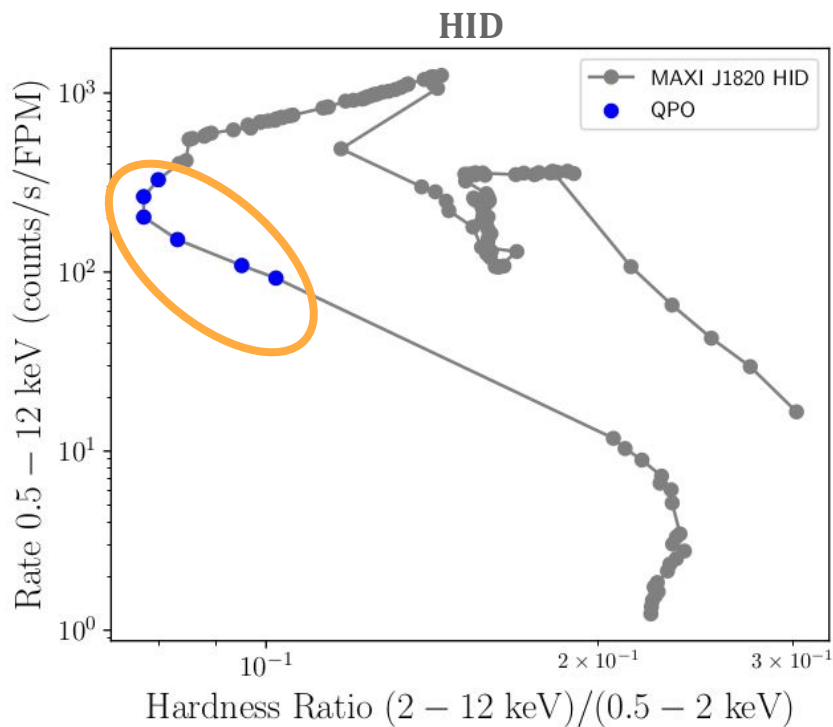
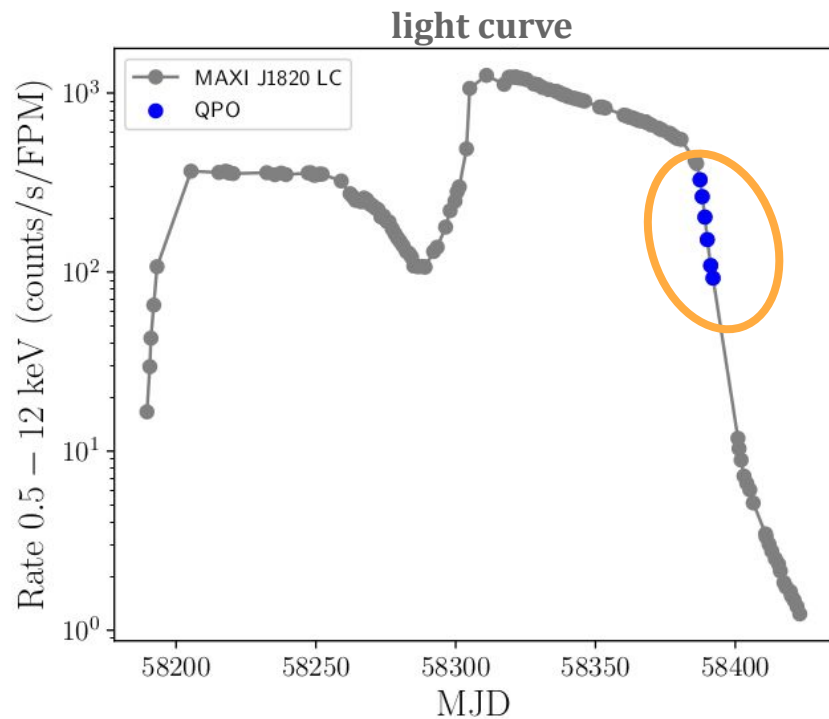
MAXI J1820+070: spectral states



MAXI J1820+070: quasi-periodic oscillations



MAXI J1820+070: the soft-to-hard transition



Fourier timing analysis

$S(\nu); H(\nu)$ FFTs of soft and hard energy light curves

Phase lags $\arg[\langle S^*(\nu)H(\nu) \rangle]$ phase of the average cross spectrum

time lag = $\frac{\text{phase lags}}{2\pi\nu}$ measures the time delay between two signals

Coherence function measures degree of linear correlation between two signals

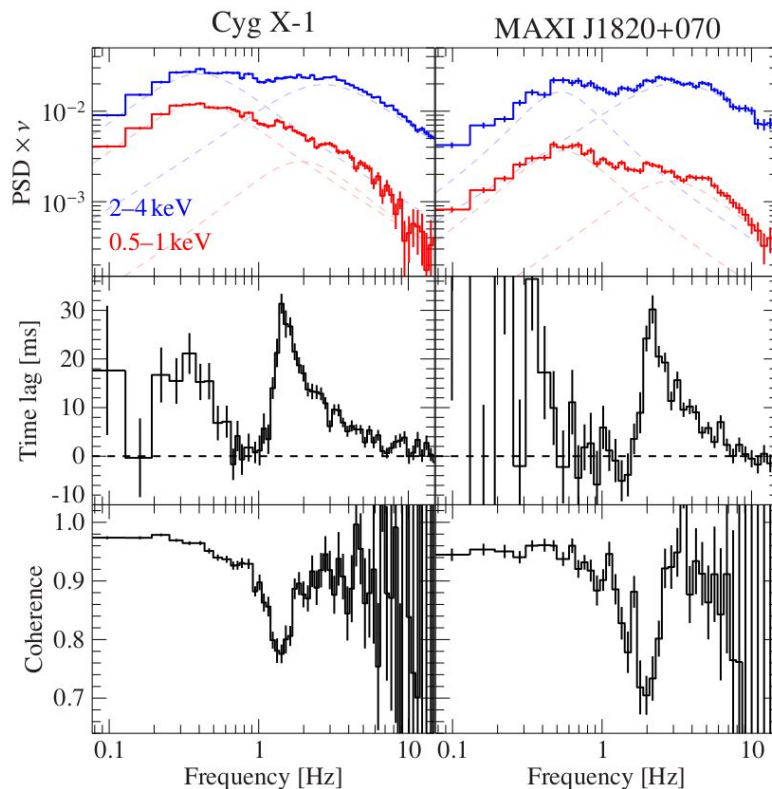
$$\gamma^2(\nu) = \frac{|\langle S^*(\nu)H(\nu) \rangle|^2}{|\langle S(\nu) \rangle|^2 |\langle H(\nu) \rangle|^2} \quad S(\nu); H(\nu) \text{ related by linear transform} \rightarrow \gamma^2(\nu) = 1$$

$\gamma^2(\nu)$ **drops** if two components with different amplitudes and phases of their cross vectors contribute to the variability over the same frequency range

Vaughan & Nowak (1997)

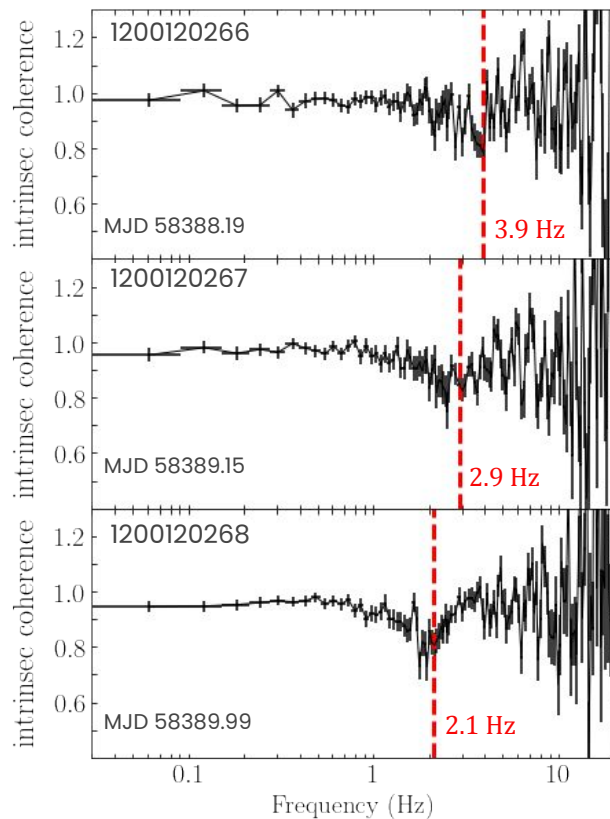
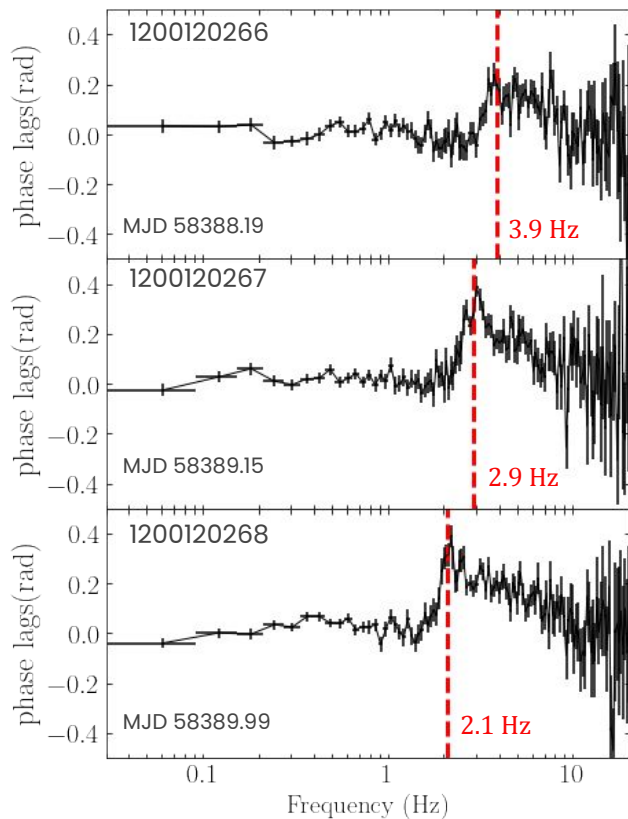
Motivation of our work

König et al 2024

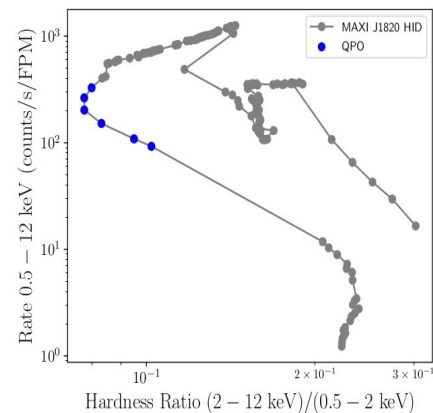


abrupt lag change
&
narrow drop in
coherence

Phase lags & coherence function

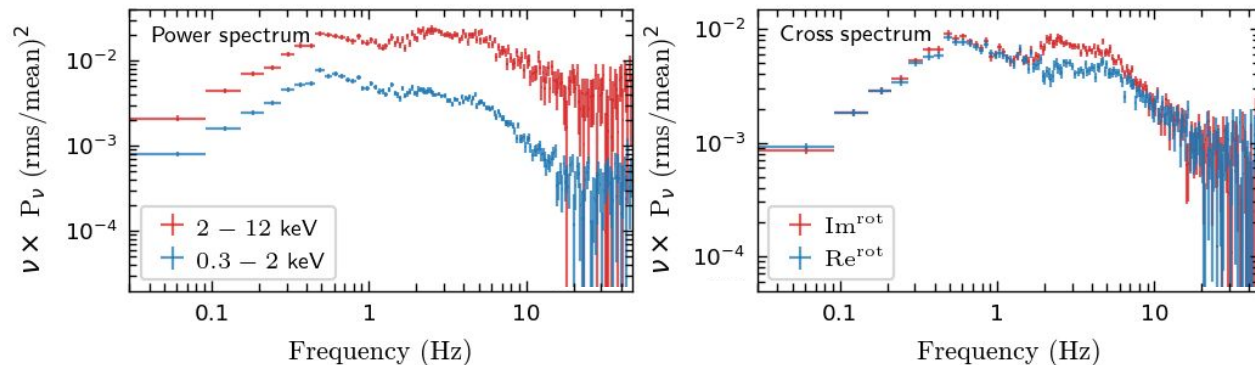


soft: 0.3 – 2 keV
hard: 2 – 12 keV



only seen when
soft band is
below 2 keV

FITTING THE PS & CS



We fit a **multi-Lorentzian model** (Mendez et al 2024)

same combination of Lorentzians
coherent in different energy bands but incoherent with each other

$$PS_{2-12} = \sum_{i=1}^n A_i L(\nu, \nu_{0,i}, \Delta_i)$$

$$\text{Im}[CS] = \sum_{i=1}^n \sqrt{A_i B_i} L(\nu, \nu_{0,i}, \Delta_i) \sin(\Delta\phi_i)$$

$$PS_{0.3-2} = \sum_{i=1}^n B_i L(\nu, \nu_{0,i}, \Delta_i)$$

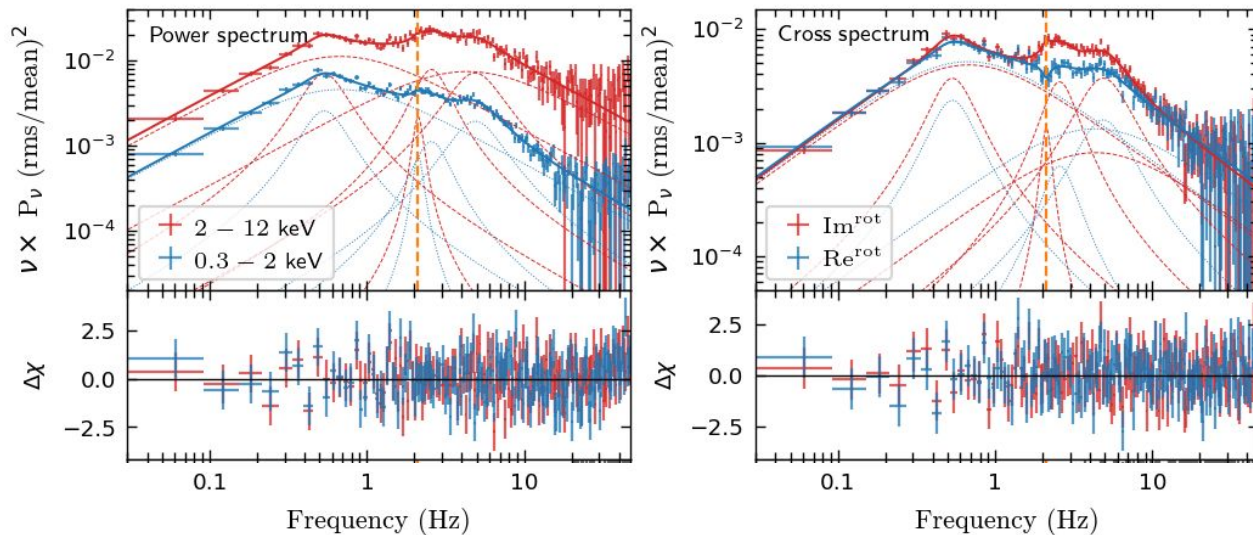
$$\text{Re}[CS] = \sum_{i=1}^n \sqrt{A_i B_i} L(\nu, \nu_{0,i}, \Delta_i) \cos(\Delta\phi_i)$$

from the best-fitting model, we **derive** the model for the lags & coherence (Mendez et al 2024)

FITTING THE PS & CS

$\chi^2 \approx 582$ for 564 d.o.f

obsID 1200120268

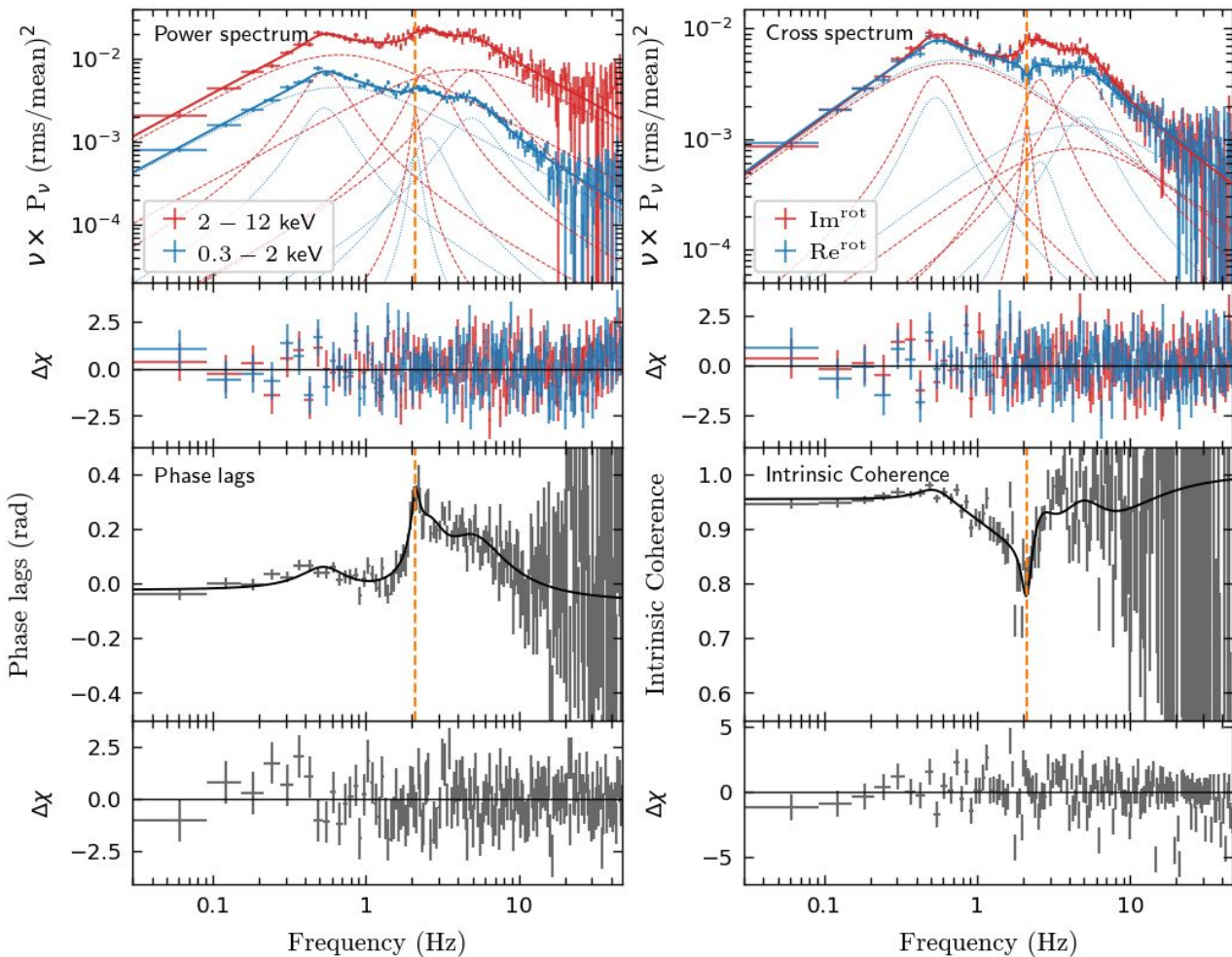


component significant only in the CS, not in the PS

small real part but large imaginary part \longrightarrow **imaginary QPO**

FITTING THE PS & CS

the imaginary QPO
explains the narrow
features



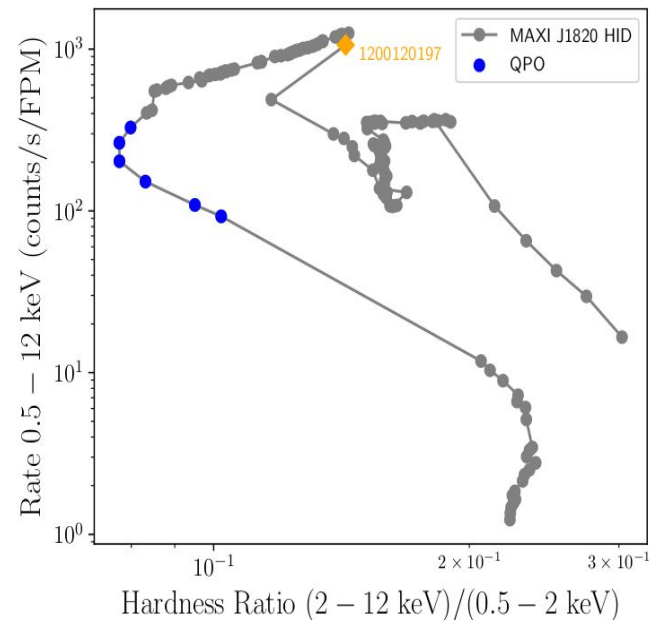
PROPERTIES OF THE IMAGINARY QPO

ΔT (days)*	ν_{QPO} (Hz)
0	6.1
0.8	6.7
1.1	3.7
1.7	2.9
2.6	2.1
3.8	1.5
4.6	1.2

* time since the first obsID considered

QPO	ν_{QPO} (Hz)	$Q = (\nu_{\text{QPO}}/\text{FWHM})$
imaginary	$\sim 1 - 7$	$\sim 2 - 14$
type B	$\sim 5 - 6$	$\gtrsim 6$
type C	$\sim 0.1 - 15$	$\sim 7 - 12$

Casella et al 2005;
Motta 2016

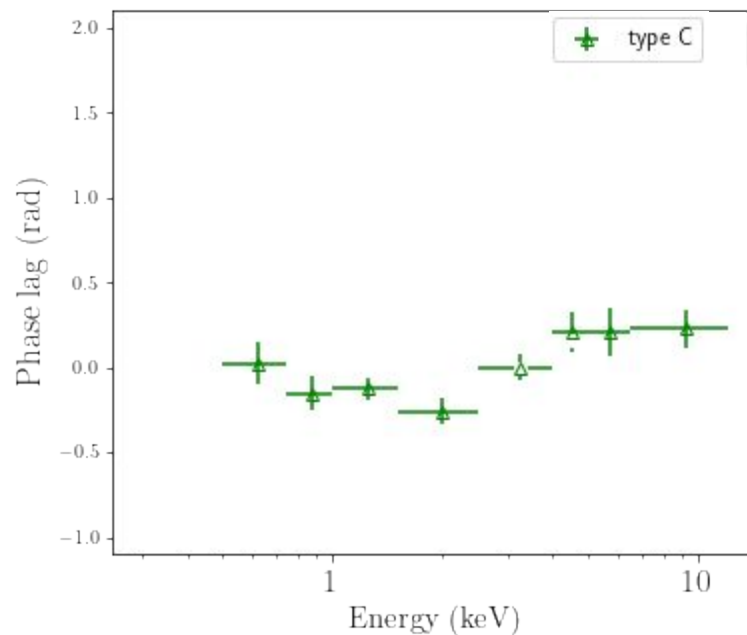
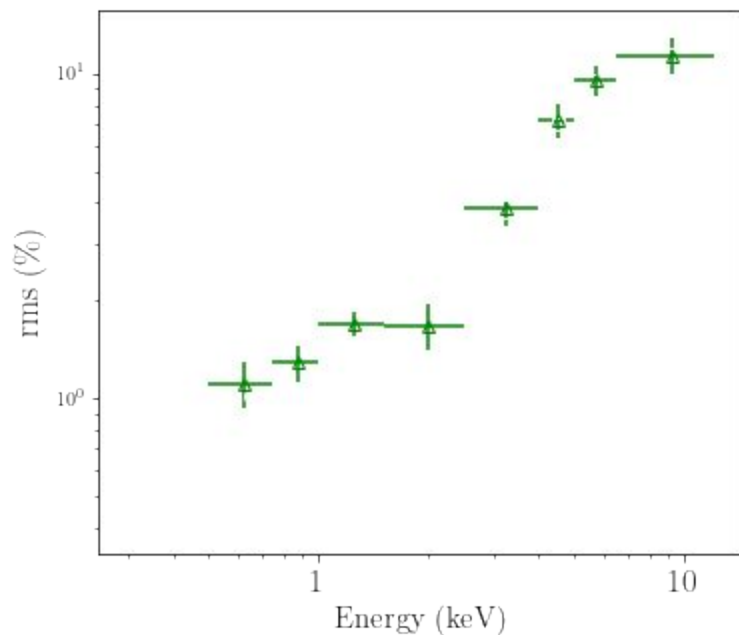


previous type C with Q between 0.2 & 7
(Alabarta et al 2022 ; Ma et al 2023)

rms & phase-lag spectra

type-C QPO in MAXI J1820+070
obsID 1200120197

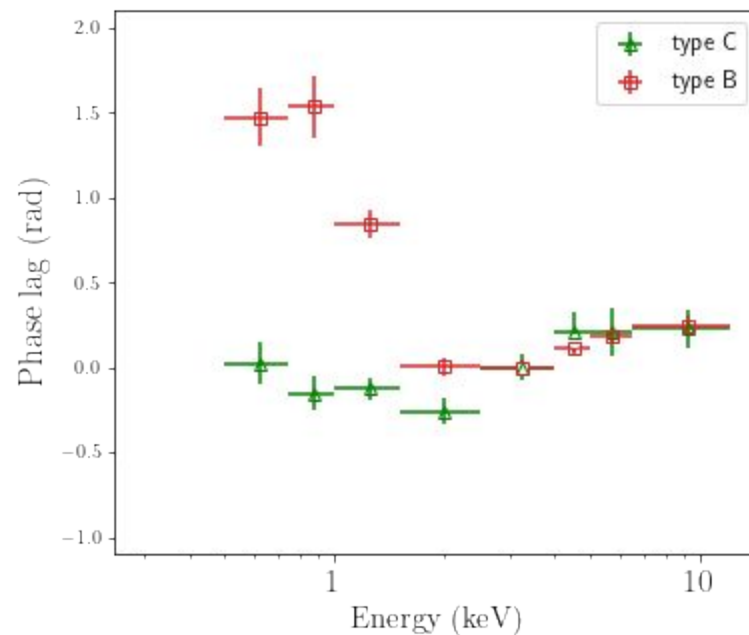
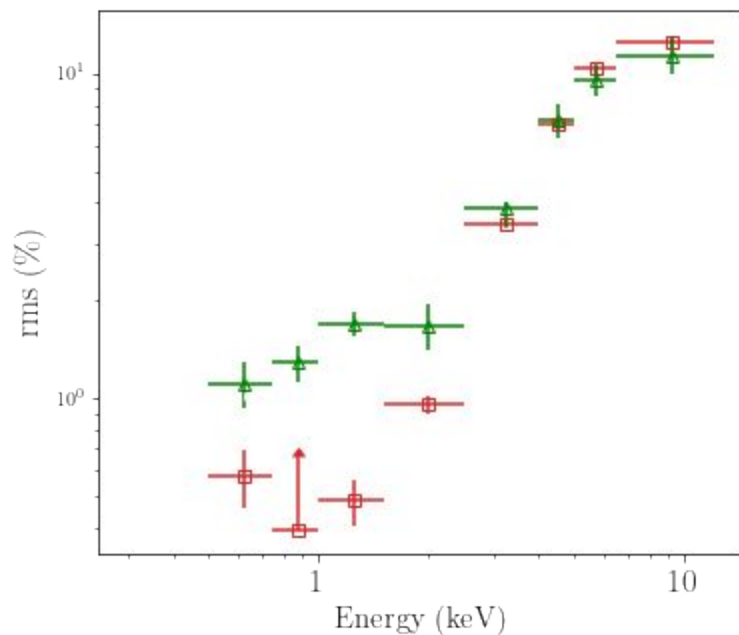
$$\nu_C = 4.4 \text{ Hz}$$



rms & phase-lag spectra

type-B QPO in MAXI J1820+070
obsID 1200120197

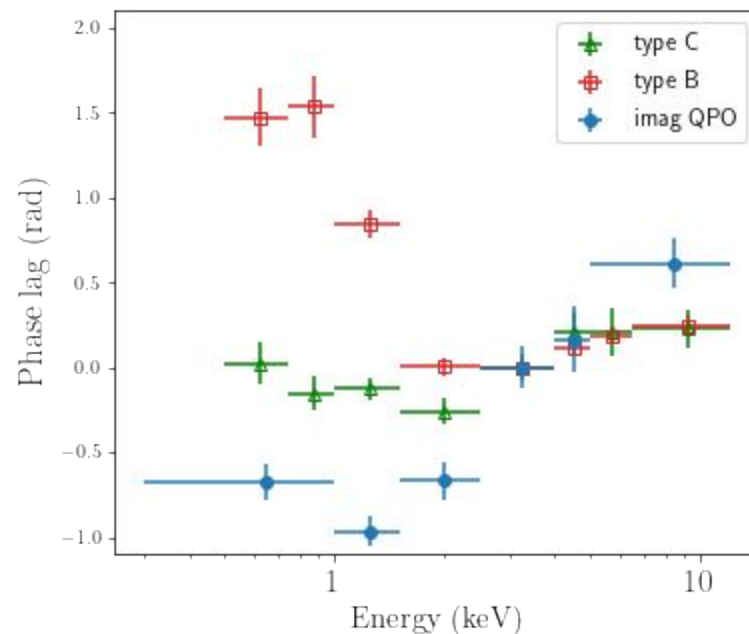
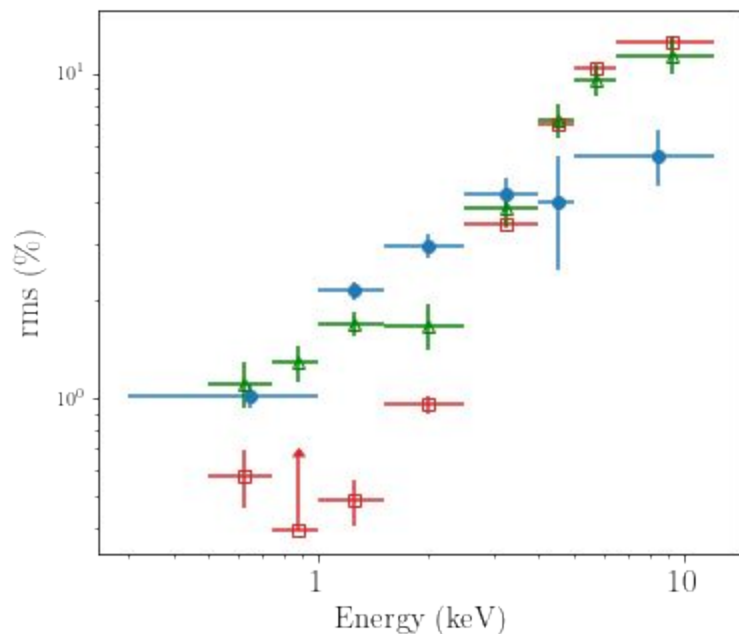
$\nu_C = 4.4$ Hz $\nu_B = 4.0$ Hz



rms & phase-lag spectra

imag QPO in MAXI J1820+070
obsID 1200120266

$$\nu_C = 4.4 \text{ Hz} \quad \nu_B = 4.0 \text{ Hz} \quad \nu_{\text{imag}} = 3.8 \text{ Hz}$$



Conclusions

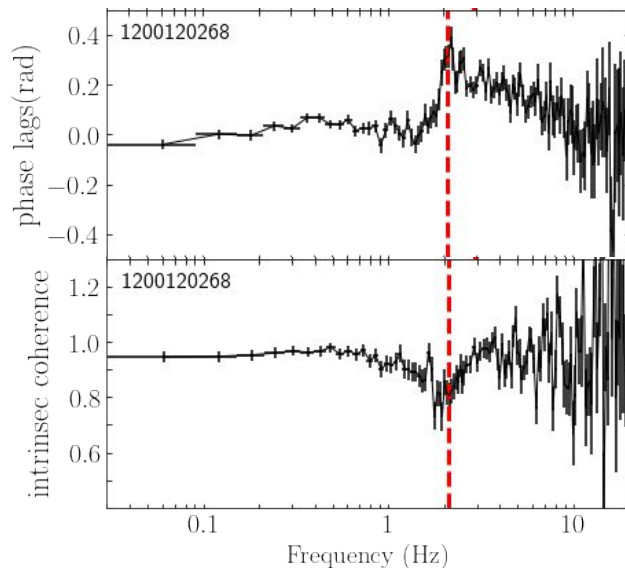
Discovery of **imaginary QPOs** along the soft-to-hard transition of MAXI J1820+070

Multi-Lorentzian model (Mendez et al 2024) fit explains the narrow drop of the coherence

Soft band needed \rightarrow component variable at low energies

Seen in the transition \rightarrow increase of inner disc radius or reappearance of corona over disc (feedback)

Study of properties of imaginary QPO, **new type of QPO?**



Future work

Analysis of sources during the decay considering soft X-ray data \rightarrow NICER data

Appendix

BEST-FITTING VALUES OF THE IMAGINARY QPO

obsID	ν_{QPO} (Hz)	FWHM (Hz)	$Q = (\nu_0/\text{FWHM})$	rms ^a _{QPO} (%)	rms ^b _{0.1-50 Hz} (%)
1200120265	6.10 ± 0.08	2.9 ± 0.3	2.1 ± 0.2	7.9 ± 0.4	18.6 ± 0.2
1200120266_p1	6.67 ± 0.09	1.2 ± 0.3	5.5 ± 1.2	4.2 ± 0.7	20.0 ± 0.2
1200120266_p2	3.75 ± 0.03	0.6 ± 0.1	6.6 ± 1.4	3.9 ± 0.8	24.4 ± 0.2
1200120267	2.96 ± 0.05	0.7 ± 0.2	4.1 ± 1.2	3.4 ± 1.1	26.5 ± 0.2
1200120268	2.11 ± 0.03	0.3 ± 0.1	6.1 ± 2.5	2.4 ± 0.8	27.4 ± 0.2
1200120269	1.55 ± 0.04	0.5 ± 0.1	3.2 ± 0.8	2.3 ± 2.0	27.0 ± 0.8
1200120270	1.26 ± 0.02	0.09 ± 0.03	14.6 ± 5.0	4.4 ± 1.1	26.7 ± 0.3
Type B	$\sim 5 - 6$		≥ 6		
Type C	$\sim 0.1 - 15$		$\sim 7 - 12$		

^a The rms amplitude of the QPO in the 2.0 – 12.0 keV band.

^b The rms amplitude of the broadband in the 0.3 – 12.0 keV band.

previous type C with Q between 0.2 & 7 (Alabarta et al 2022 ; Ma et al 2023)

Need of narrow component

